Using astrophysical knowledge in gravitational-wave data analysis of binary inspirals

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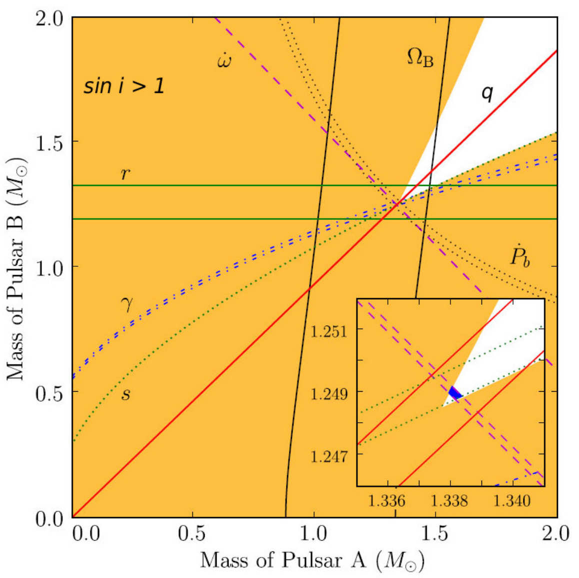
Outline

Gravitational waves

GWs:

- "Ripples in spacetime"
- Predicted by Einstein's theory of General Relativity
- *Indirectly* observed for the Hulse-Taylor binary pulsar:

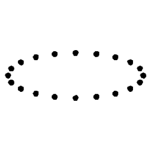


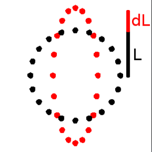
(Breton et al., Science, 2008)

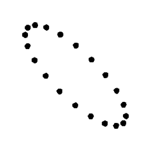
Gravitational waves

Gravitational waves. . .

- propagate transversely at the speed of light
- are quadrupole radiation at the lowest order
- stretch and squeeze spacetime in two polarisations
- allow us to measure their amplitude

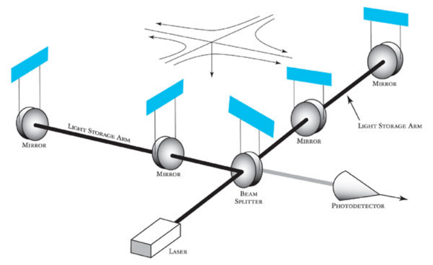


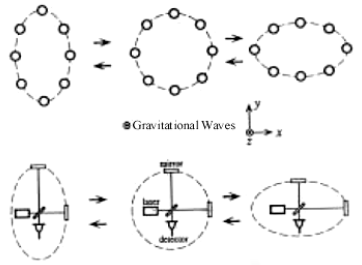




- Strain: $h(t) = h_+(t)F_+(t) + h_\times(t)F_\times(t) = \frac{\delta L(t)}{L} \sim 10^{-22}$

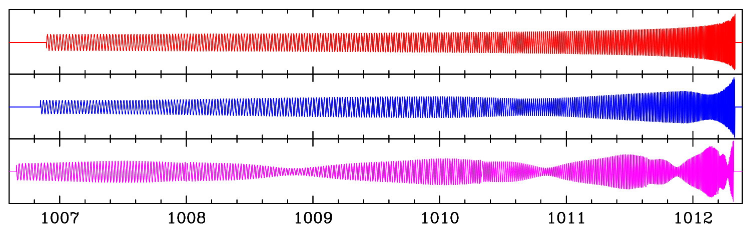
Laser Interferometer GW Observatory (LIGO)

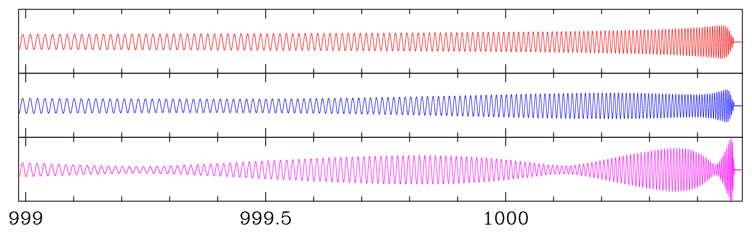




Inspiral waveforms with increasing spin

LIGO and Virgo detect the last ~ 10 s of a binary inspiral:



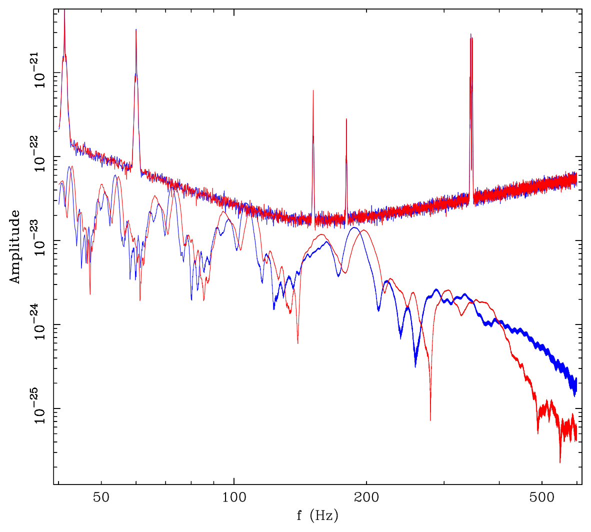


$a_{\rm spin} \equiv S/M^2 = 0.0$, 0.1 and 0.5

Signal injection into detector noise

Example:

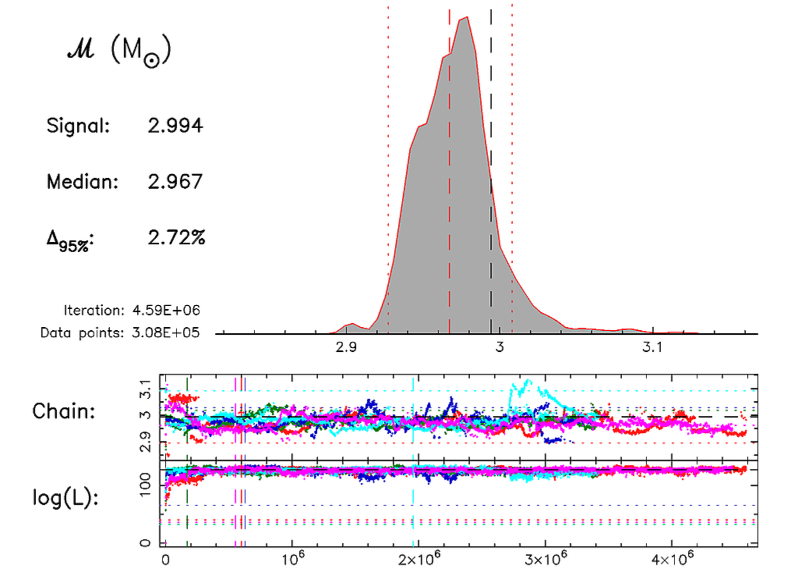
- Using two 4-km detectors H1, L1
- Inject signal coherently
- Σ SNR = 17
- Retrieve physical parameters using MCMC



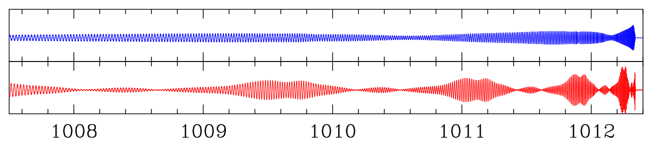
SPINSPIRAL code

- Use Markov-Chain Monte Carlo for parameter estimation
- Follow-up after detection
- Gaussian, stationary noise or LIGO/Virgo detector data
- Analyse software injections, hardware injections, detection candidates/interesting events
- Include spin in injections and analysis
- Use any network composed of LIGO/Virgo detectors:
  - $\mathrm{PDF}(\vec{\lambda}) \propto \mathrm{prior}(\vec{\lambda}) \times \prod_i L_i(d|\vec{\lambda})$
- Result: posterior probability-density function (PDF) of the parameter set that describes the model (9–12–15 D)

SPINSPIRAL example



Correlations increase with spin

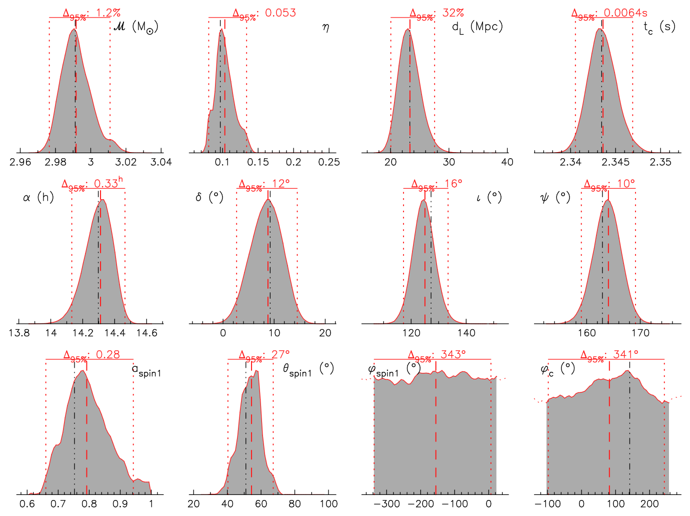


	M_c	η	a_{spin}	ϑ_{SL}	R.A.	Dec.
M_c		0.22	0.42	0.17	−0.40	0.19
η	−0.27		−0.34	−0.53	−0.07	−0.04
a_{spin}	−0.61	0.89		−0.04	0.11	0.62
ϑ_{SL}	0.66	−0.87	−0.99		0.02	−0.34
R.A.	−0.36	0.01	0.02	−0.02		0.12
Dec.	−0.23	0.08	0.18	−0.20	−0.05	

Parameters:

- BH-NS
- H1 & L1
- $M_1 = 10\,M_\odot$
- $M_2 = 1.4\,M_\odot$
- $a_{spin} = 0.1, 0.8$
- $\theta_{SL} = 55°$
- Network SNR ≈ 25

MCMC results for the analysis of a BH-NS signal

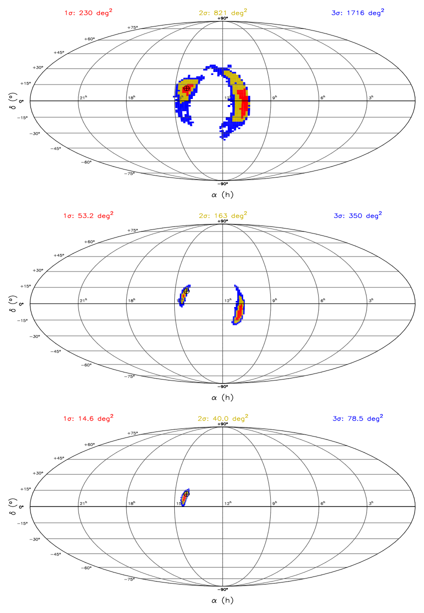


van der Sluys et al., 2008

Parameters:

- H1, L1, V
- $M = 10, 1.4\, M_\odot$
- $d_L = 22.4$ Mpc
- $a_{\rm spin} = 0.8$, $\theta_{\rm SL} = 55°$
- Σ SNR ≈ 17.0
- simulated noise
- Black dash-dotted line: injection
- Red dashed line: median
- Δ's: 95% probability

Sky position for signals with different spins



Spinning BH, non-spinning NS:

$10 + 1.4\,M_\odot$, 16–22 Mpc, Σ SNR=17

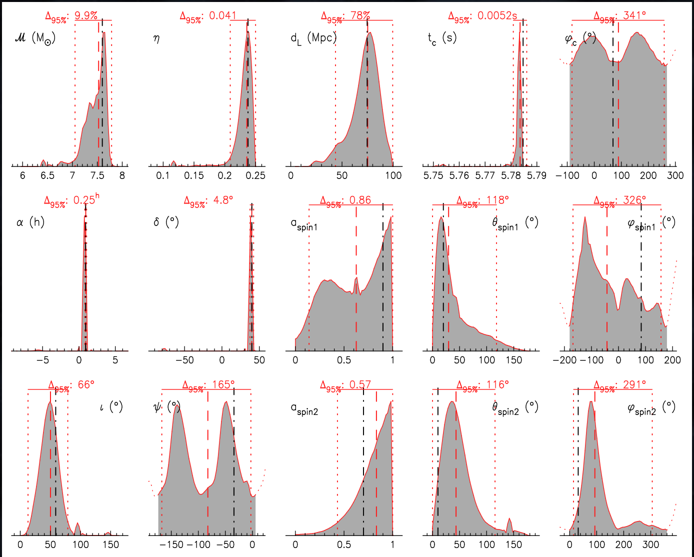
2 detectors, $a_{\rm spin} = 0.0$
2-σ accuracy: $821^{\circ 2}$

2 detectors, $a_{\rm spin} = 0.5$
2-σ accuracy: $163^{\circ 2}$

3 detectors, $a_{\rm spin} = 0.5$
2-σ accuracy: $40^{\circ 2}$

van der Sluys et al., 2008; Raymond et al., 2009; **Poster by Ben Farr**

Analysis of a BH-BH signal with spins

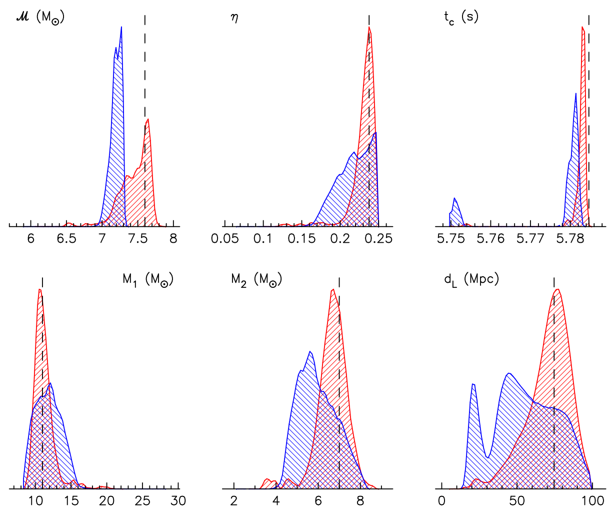


van der Sluys et al., in preparation

HS-2:

- 3.5-pN waveform
- 3 detectors (H1,L1,V)
- $\mathcal{M} = 7.6\,M_\odot$, $\eta = 0.238$; $M_1 = 11.0\,M_\odot$, $M_2 = 7.0\,M_\odot$
- $a_{s1,2} = 0.9, 0.7$
- $\theta_{s1,2} = 10, 20°$
- $d_L = 74.5$ Mpc
- Σ SNR=15
- simulated noise

The importance of having spins in your analysis



Signal with spins

Analysis with spinning template

Analysis with non-spinning template

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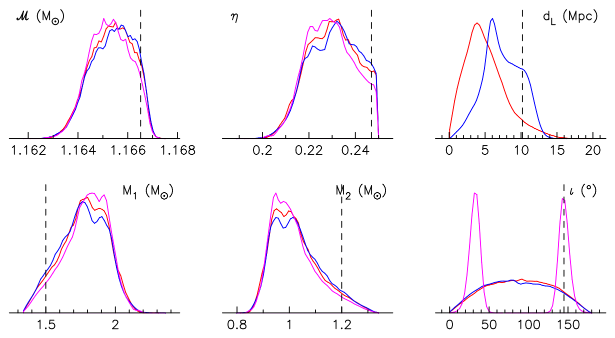
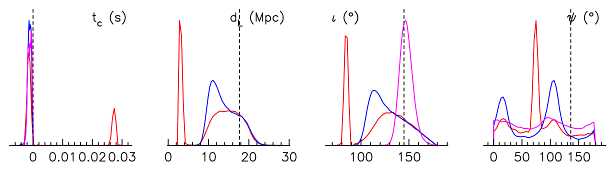
Using astrophysical data to constrain parameters

1 detector (H1):



3 detectors (H1,L1,V):



NS-NS, non-spinning:
$1.2 + 1.5\, M_\odot$
$d_L \approx 10.2 - 17.8$ Mpc
(Σ SNR=15.0)

No astrophysical information

Sky position known

Sky position and distance known

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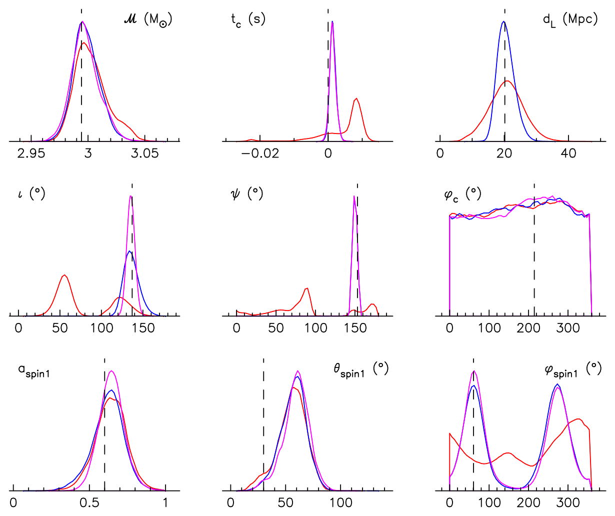
See also: **Nissanke et al., 2010**

Using astrophysical data to constrain parameters

2 detectors (H1,L1):



BH-NS, spinning BH:
$10. + 1.4\, M_\odot$
$d_L \approx 20.2$ Mpc
(Σ SNR=15.0)

No astrophysical information

Sky position known

Sky position and distance known

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Conclusions

SPINSPIRAL

- SPINSPIRAL can recover the 12–15 parameters of a binary inspiral, including one or two spins, using an MCMC technique

- Sky-position reconstruction (few $\times 10^{\circ^2}$) is poor for astrophysical standards
- Combination of position, distance and time can lead to association with an electromagnetic detection (*e.g.* GRB)

Taking into account spins

- The inclusion of spin adds significantly to the number of dimensions (9–12–15) and introduces (strong) correlations
- Failing to take into account spin can result to biases in especially mass parameters

Conclusions (numbers are preliminary)

Using astrophysical knowledge for GW data analysis: no spins

- Knowing the sky position of a source improves determination of:
  - distance ($\sim 20-50\%$)
  - inclination
- Knowing the position *and distance* improves inclination further, also in 1-detector analysis

Using astrophysical knowledge for GW data analysis: spins

- Knowing the sky position of a source improves determination of:
  - distance ($\sim 50\%$)
  - inclination, polarisation angle ($50-90\%$)
  - masses ($\sim 20\%$)
  - spin angles
- Knowing the position *and distance* improves:
  - spin magnitude ($\sim 20\%$)

End. . .

Predicted detection rates

Realistic estimate:

	Rates (yr^{-1})			Horizon (Mpc)		
	NS-NS	BH-NS	BH-BH	NS-NS	BH-NS	BH-BH
Initial	0.015	0.004	0.01	32	67	160
Enhanced	0.15	0.04	0.11	71	149	349
Advanced	20	5.7	16	364	767	1850

Plausible, optimistic estimate:

	Rates (yr^{-1})			Horizon (Mpc)		
	NS-NS	BH-NS	BH-BH	NS-NS	BH-NS	BH-BH
Initial	0.15	0.13	1.7	32	67	160
Enhanced	1.5	1.4	18	71	149	349
Advanced	200	190	2700	364	767	1850

Estimates assume $M_{\rm NS} = 1.4\, M_\odot$ and $M_{\rm BH} = 10\, M_\odot$

CBC group, rates document

MCMC analyses

MCMC parameters

Masses: $\mathcal{M} \equiv (M_1 + M_2)\,\eta^{3/5}$ & $\eta \equiv \frac{M_1 M_2}{(M_1+M_2)^2}$, distance: $\log d_L$, time and phase at coalescence: t_c & φ_c, position: α & $\sin\delta$, spin magnitude: $a_{\mathrm{spin}_{1,2}}$, spin orientation: $\cos\theta_{\mathrm{spin}_{1,2}}$ & $\varphi_{\mathrm{spin}_{1,2}}$ & binary orientation: $\cos(\iota)$ & ψ

MCMC set-up

- ≥ 5 serial chains per run, starting from offset parameter values
- Chain length: $\sim$ few $\times 10^6$ states; burn-in: $\sim$ few $\times 10^5$ states
- Run time: 10 days on a 2.8 GHz CPU for 1.5-pN waveform; $\sim 2.5\times$ longer for 3.5-pN

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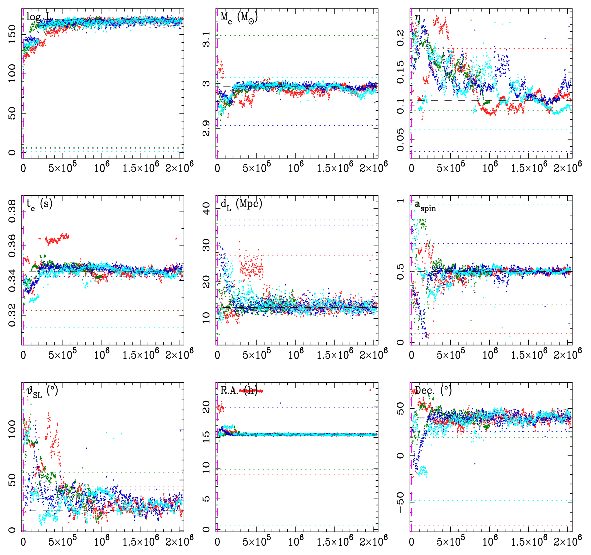
Analysis details: BH-NS signal

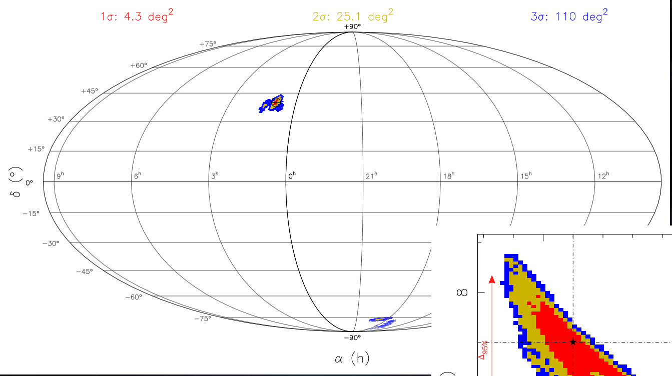
- Signals injected in simulated noise for H1L1V @ SNR $\approx$17.0
- Fiducial binary: $M_{1,2} = 10 + 1.4\,M_\odot$, $d_L = 16$–23 Mpc
- Spin: $a_{\mathrm{spin}} = 0.0, 0.1, 0.5, 0.8$, $\theta_{\mathrm{SL}} = 20^\circ, 55^\circ$

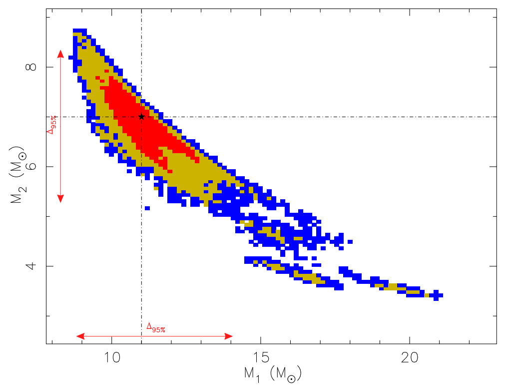
Convergence of chains



- Dots: starting values
- Dashes: injection values

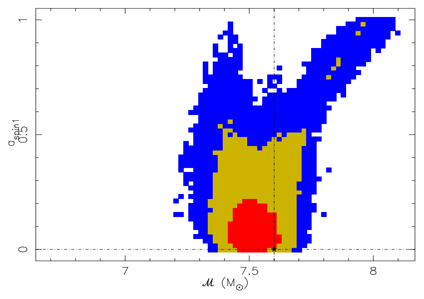
Analysis of a BH-BH signal with spins



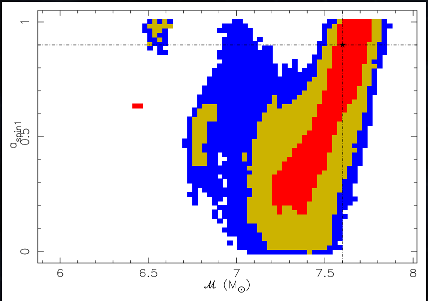


van der Sluys et al., in preparation

The nuisance of having spins in your analysis



Signal **without** spins,
analysis with spinning template



Signal **with** spins,
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